## DIRECT OBSERVATION OF EXOPLANETS WITH GRAVITY

VLTI SCHOOL 2021 S. LACOUR

# ON THE DIFFICULTY OF DIRECT DETECTION











Bern model of planet formation and evolution (Mordasini et al. 2012) Including the Young Stellar formation regions (Cha, Sco-Cen) Including Young moving groups (Beta Pictoris, TW Hya, Carina, etc..)



#### Stellar light (Extreme AO, 95%Strehl)







500 mas

EC=58/60%;t0=8.1/9.7ms;texp=420/420s





Fomalhaut b Planet 2006 17/06/2021 9



17/06/2021

# THE GRAVITY INTERFEROMETER





## GRAVITY @ ESO VLTI





The GRAVITY cryostat



Haug, Thiel, Hausmann & the GRAVITY collaboration











Haug, Thiel, Hausmann & the GRAVITY collaboration

## FIBERCOUPLER



## FIBERCOUPLER







 $\delta OPD = \vec{B} \cdot (\vec{\alpha} - \vec{\beta})$ 



	Metrology phases	Fringe Tracker visibilities
	O  O O	
	V.14      V0.1      2      3      4      6      6      7      8      9      10      11      6        1      40.565004      40.564764      40.5195026      40.5195626	
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#### Science camera visibilities

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![](_page_32_Figure_0.jpeg)

# RESULTS FROM OPTICAL INTERFEROMETRY

![](_page_33_Picture_1.jpeg)

![](_page_34_Figure_1.jpeg)

![](_page_35_Figure_0.jpeg)

![](_page_35_Figure_1.jpeg)

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36

Nowak et al. (2020) Lagrange et al. (2020) → Press release tomorrow

![](_page_36_Figure_2.jpeg)

![](_page_37_Figure_0.jpeg)

![](_page_37_Figure_1.jpeg)

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![](_page_38_Figure_1.jpeg)

![](_page_39_Figure_1.jpeg)

![](_page_40_Figure_1.jpeg)

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**4**T

![](_page_41_Figure_1.jpeg)

![](_page_42_Figure_0.jpeg)

5
2

Parameter	Prior distribution	Poster	Unit		
Star		βF			
Stellar mass	Gaussian $(1.77 \pm 0.02)$	1.82	$M_{\odot}$		
Parallax	Uniform	51.42	$\max$		
proper motion	Uniform	$\texttt{ra}{=}4.88\pm0.02$	m mas/yr		
$RV v_0$	Uniform	-23.0	m/s		
RV jitter	Uniform	31.2	$1\pm 8.09$	m/s	
Planets		$\beta$ Pictoris b	$\beta$ Pictoris c		
Semi-major axis	Log Uniform	$9.90\pm0.05$	$2.72\pm0.02$	au	
Eccentricity	Uniform	$0.10\pm0.01$	$0.37\pm0.12$		
Inclination	Sin Uniform	$88.99 \pm 0.01$	$89.17 \pm 0.50$	$\deg$	
PA of ascending node	Uniform	$31.82\pm0.02$	$30.98 \pm 0.12$	$\deg$	
Argument of periastron	ent of periastron Uniform		$196.9 \pm 3.5$ $66.2 \pm 2.5$		
Epoch of periastron	ch of periastron Uniform		$0.83\pm0.02$		
Planet mass	Gaussian $(15.4 \pm 3.0)$		$9.0 \pm 1.6$		
	Log Uniform		$8.2\pm0.8$	$M_{jup}$	
Period		$23.28\pm0.46$	$3.37\pm0.04$	years	
$\Delta m_{\rm K}$		$8.9 \pm 0.1$	$10.8\pm0.1$	mag	

Nowak et al. (2020)

## PDS 70 B & C

## PDS 70 B & C

Gaussian prior on coplanarity with disk : Sigma = 10 degrees

Gaussian prior on coplanarity between b & c

Wang et al. (work in progress)

![](_page_44_Figure_4.jpeg)

![](_page_45_Figure_0.jpeg)

2000 Rjup

 $\rightarrow$ 

![](_page_46_Figure_0.jpeg)

With circumplanetary model, for PDS70b: Rcpd < 1mas (0.1au)

> R Hill = 1.6 au R Bondi = 8 au

![](_page_47_Figure_0.jpeg)

### K BAND SPECTRA (J. WANG ET AL. IN PREPARATION)

![](_page_48_Figure_0.jpeg)

![](_page_48_Figure_1.jpeg)

- PLANETARY-TYPE ATMOSPHERE
- NO BLACK BODY CPD
- NEED TO ADD EXTINCTION

## BETA PICTORIS B

- 23 MYR OLD
- 20 PC
- ~ 12 MJUP PLANET

![](_page_49_Picture_4.jpeg)

### Beta Pictoris b

![](_page_50_Figure_1.jpeg)

#### Gravity Collaboration: M. Nowak et al. (2020)

### Spectral Modeling

#### Self-consistent Chemical Modeling Cloud Modeling Radiative-convective temperature => "correct" / slow / can use sparse data

#### Free retrieval

Abundance Cloud opacity Temperature Parametrize! => "data speak" / fast / compare models

- complicated
- unknown
- love to blow up self-consistent codes

![](_page_52_Figure_0.jpeg)

![](_page_53_Figure_0.jpeg)

### Planet composition varies with disk radius

![](_page_54_Figure_1.jpeg)

### This simple picture ignores chemistry (freeze out only)

![](_page_55_Figure_0.jpeg)

![](_page_55_Figure_1.jpeg)

# PROSPECTS FOR OPTICAL INTERFEROMETRY

![](_page_56_Picture_1.jpeg)

![](_page_57_Figure_1.jpeg)

8

![](_page_58_Figure_0.jpeg)

![](_page_59_Figure_0.jpeg)

The Very Large Telescope in 2030 – White paper proposed to ESO – strongly supported

## GRAVITY<sup>+</sup> : Towards foint science, all sky milliarcsec optical interferometric imaging

### STC recommends immediate start

#### GRAVITY<sup>+</sup>: Towards faint science, all sky milliarcsecond optical interferometric imaging

Constructions by Performing and the Construction of the Constructi

#### Introduction

Ever since the development of 'aperture synthesis' spatial interferometry in the raband in the 1950s (Syr& Hewshi 1946), this technique has been the standard choice for development of large aperture, high angular resolution tolescopes in the microwave a radio bands (Thompson, Moran & Swenson 1986).

The probability of radia interference type has based on assuming the range has much shifting the same based on the same set of the same set of the same set of the same set of the same based on the same set of the same based on the same set of the same s

> s over a wide field-of-view. evel polarimetric m extral astrometry at the level of 0.06 degrees rms in interfer

![](_page_60_Picture_10.jpeg)

### **Off Axis Tracking**

### Adaptive Optics

### Laser Guide Stars

**Improved Sensitivity** 

## ExoGRAVITY : exoplanets are design-driver science case

Credit: ESO, Huedepohl