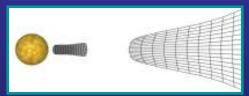
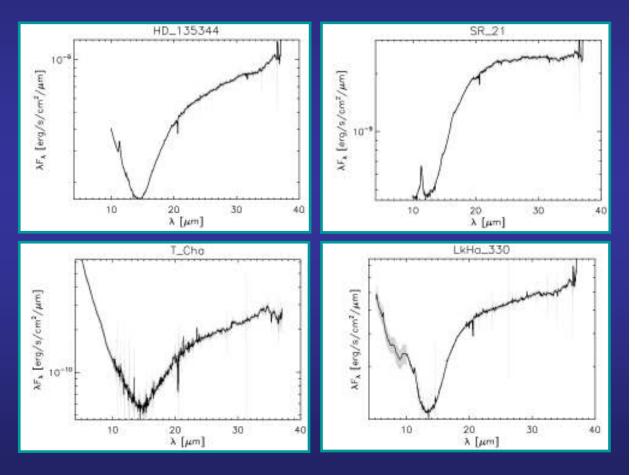
Probing the inner disk and gap

- Scientific background :
- Spitzer/IRS observations (5-35 µm)
- 4 sources selected on their spectra
 - Optically thin to optically thick disk (see Alexander's talk)
 - Possible explanation : a large gap in the disk



- Question: can a disk model reproduce the SED ?

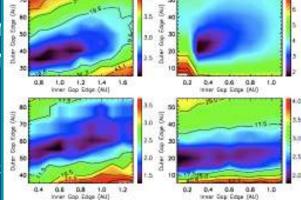


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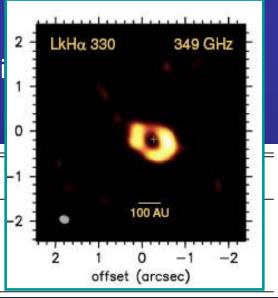
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Scientific background

- Modeling done by Brown et al. (2007), using 2D radiative transfer code RADMC (Dullemond & Dominik 2004)
- χ^2 minimization process on the SED of each source
 - Derivation of the parameters of the possible gaps (R_{min}, R_{max} and inclination)
- For LkHα 330, confi observations (Brown
- Does interferometry (R<Gap_{inner}) and thu



limeter



Source	Spectral Type	A_v (mag)	Distance (pc)	$_{\rm (M_{\odot})}^{\rm M_{star}}$	$_{\rm (K)}^{\rm T_{\rm eff}}$	$_{\rm (R_{\odot})}^{\rm R_{star}}$	${\scriptstyle ({ m AU})}^{ m Gap_{inner}}$	${\scriptstyle { (AU) } \atop (AU) }$
LkHa 330 T Cha	G3 ^a G8 ^c	$\frac{1.8^{a}}{1.5^{c}}$	250 ^b 66 ^d	$2.5 \\ 1.5$	6200 5900	$3.3 \\ 1.1$	$0.8 \\ 0.15$	35 20
HD 135344 SR 21	F4 ^e G2.5 ^g	0.47 ^f 9 ^g	$\frac{84}{160}$	$1.8 \\ 2.5$	6600 6300	$1.9 \\ 2.6$	$0.55 \\ 0.35$	$55 \\ 20$

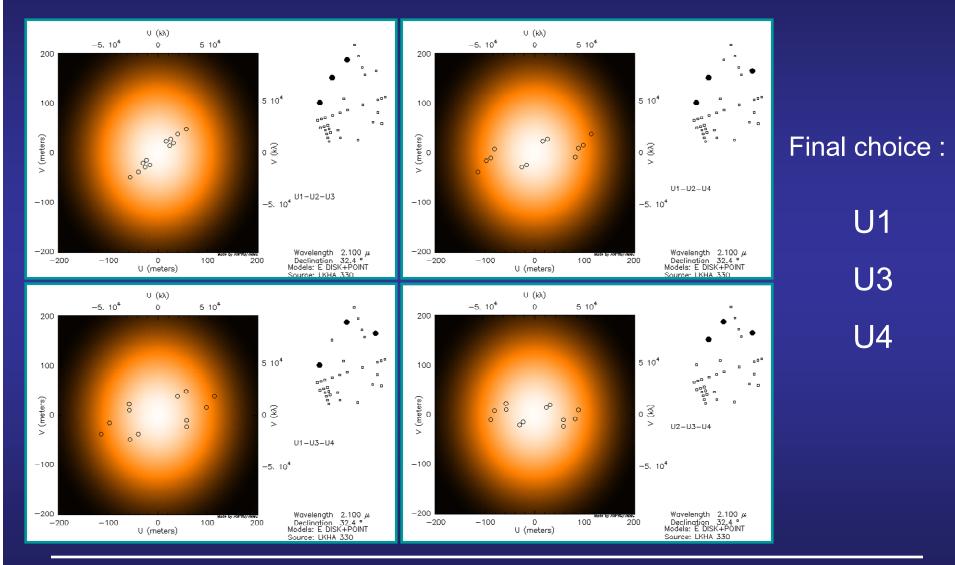
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How?

- AMBER with UTs : why ?
 - Direct imaging is not possible
 - Sources too faint for MIDI and for AMBER with ATs
- \Rightarrow AMBER with UTs is the only instrument at VLTI for these observations
- Details :
 - Observable from Paranal
 - Magn K : 5.84 6.9, Magn V : 8.6 14.2 (<17)
 - Low resolution (magnitude limit)
 - 3 hours of observation (H.A.: -2 to +1)
 - Model used :
 - Point = star (flux : 0.7)
 - E-disk = inclined disk to reproduce the parameters inferred from the modeling (inclination known, flux : 0.3)
 - Calibrators found for each stars

Which baselines ?

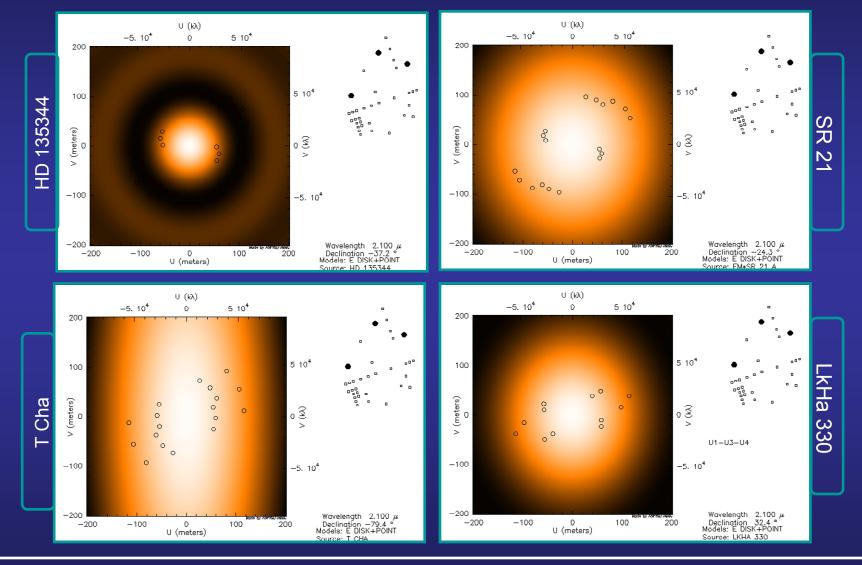


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Slide 4

UV coverage

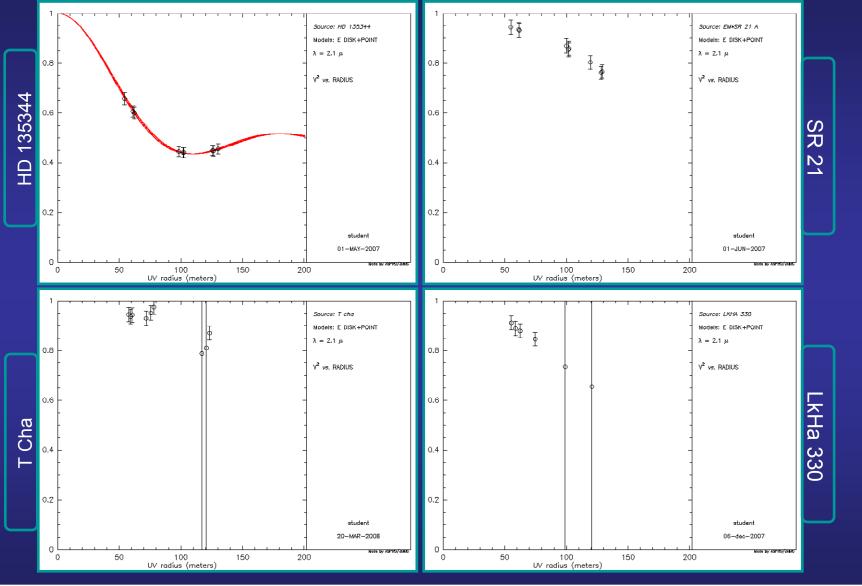


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Slide 5

Visibilities



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Slide 6